



# HE0359-3959: An Extremely Radiating Quasar

M. L. Martínez-Aldama<sup>1\*</sup>, A. Del Olmo<sup>1</sup>, P. Marziani<sup>2</sup>, C. A. Negrete<sup>3</sup>, D. Dultzin<sup>4</sup> and M. A. Martínez-Carballo<sup>1</sup>

<sup>1</sup> Instituto de Astrofísica de Andalucía, IAA-CSIC, Granada, Spain, <sup>2</sup> INAF, Osservatorio Astronomico di Padova, Padua, Italy,

<sup>3</sup> CONACYT Research Fellow, Instituto de Astronomía, Universidad Nacional Autónoma de México, México City, México,

<sup>4</sup> Instituto de Astronomía, Universidad Nacional Autónoma de México, México City, México

## OPEN ACCESS

### Edited by:

Jirong Mao,  
Yunnan Observatories, National  
Astronomical Observatories (CAS),  
China

### Reviewed by:

Andrea Marinucci,  
Roma Tre University, Italy  
Milan S. Dimitrijevic,  
Astronomical Observatory, Serbia

### \*Correspondence:

M. L. Martínez-Aldama  
maryloli@iaa.es

### Specialty section:

This article was submitted to  
Milky Way and Galaxies,  
a section of the journal  
Frontiers in Astronomy and Space  
Sciences

**Received:** 31 August 2017

**Accepted:** 09 October 2017

**Published:** 28 November 2017

### Citation:

Martínez-Aldama ML, Del Olmo A,  
Marziani P, Negrete CA, Dultzin D and  
Martínez-Carballo MA (2017)  
HE0359-3959: An Extremely  
Radiating Quasar.  
Front. Astron. Space Sci. 4:29.  
doi: 10.3389/fspas.2017.00029

We present a multiwavelength spectral study of the quasar HE0359-3959, which has been identified as an extreme radiating source at intermediate redshift ( $z = 1.5209$ ). Along the spectral range, the different ionic species give information about the substructures in the broad line region. The presence of a powerful outflow with an extreme blueshifted velocity of  $\sim -6,000 \pm 500 \text{ km s}^{-1}$  is shown in the  $\text{Civ}\lambda 1549$  emission line. A prominent blueshifted component is also associated with the  $1900\text{\AA}$  blend, resembling the one observed in  $\text{Civ}\lambda 1549$ . We detect a strong contribution of very the low-ionization lines,  $\text{FeII}$  and Near-Infrared  $\text{Ca II}$  triplet. We find that the physical conditions for the low, intermediate, and high-ionization emission lines are different, which indicate that the emission lines are emitted in different zones of the broad line region. The asymmetries shown by the profiles reveal different forces over emitter zones. The high-ionization region is strongly dominated by radiation forces, which also affect the low and intermediate-ionization emitter region, commonly governed by virial motions. These results support the idea that highly radiating sources host a slim disk.

**Keywords:** quasars: emission lines, quasars: outflows, quasars: individuals HE0359-3959, quasars: supermassive black holes, galaxy evolution: feedback

## 1. EXTREME POPULATION A SOURCES ALONG THE 4DE1 MAIN SEQUENCE

The 4D Eigenvector 1 (4DE1) parameter space offers a formalism to distinguish and classify type 1 Active Galactic Nuclei (AGN) considering their spectral properties (Sulentic et al., 2000a,b). The Full Width at Half Maximum (FWHM) of  $\text{H}\beta$  broad component ( $\text{H}\beta_{\text{BC}}$ ), the strength of optical  $\text{FeII}$  blend at  $4,570\text{\AA}$  described by the ratio  $R_{\text{FeII}} = I(\text{FeII})/I(\text{H}\beta_{\text{BC}})$ , the velocity shift of the  $\text{Civ}\lambda 1549$  profile, and soft X-ray photon index ( $\Gamma_{\text{soft}}$ ), provide four observationally independent dimensions of the Eigenvector 1. In the 4DE1 optical plane, the type 1 AGN occupy a well-defined sequence, driven mainly by the Eddington ratio,  $L/L_{\text{Edd}}$ . Along this sequence we observe a variation of the physical parameters and orientation. Then, 4DE1 could be revealing an evolution sequence for type 1 AGN (Sulentic et al., 2000a; Marziani et al., 2010; Zamfir et al., 2010). For more information about the 4DE1 and update of results, see Marziani et al. in this volume.

Using the 4DE1 we identify two populations with different spectral features: A and B. Population A has a FWHM ( $\text{H}\beta_{\text{BC}}$ )  $\leq 4,000 \text{ km s}^{-1}$ . It shows large blue asymmetries in the high-ionization lines like  $\text{Civ}\lambda 1549$ , and it is majority populated by radio quiet sources. In contrast, population B shows a FWHM ( $\text{H}\beta_{\text{BC}}$ )  $> 4,000 \text{ km s}^{-1}$  and it is mostly composed of radio-loud sources (Sulentic et al., 2002; Zamfir et al., 2010). Each population can be divided into small bins with  $\Delta\text{FWHM}$

( $H\beta_{BC}$ ) = 4,000 km s<sup>-1</sup> and  $\Delta R_{FeII} = 0.5$ , defining subpopulations shown in the **Figure 1**. In this paper we focus in the subpopulation A3 and A4 ( $R_{FeII} > 1$ ), which have been identified as highly radiating sources (xA, Marziani and Sulentic, 2014). These kind of sources show high Eddington ratios ( $L/L_{Edd} > 0.2$ ) probably produced by a slim disk, which is geometrically and optically thick and it could be formed in an advection-dominated accretion flow (Abramowicz et al., 1988; Abramowicz and Straub, 2014).

We have found selection criteria to identify the xA sources based on the 4DE1 formalism. In the optical region they show a  $R_{FeII} > 1$  (high intensity of FeII) and in the UV range  $AlIII\lambda 1860/SiIII\lambda 1892 \geq 0.5$  and  $CIII\lambda 1909/SiIII\lambda 1892 \leq 1.0$  (Marziani and Sulentic, 2014). Also, they show strong blueshifted components associated with the high ionization lines, for example in  $CIV\lambda 1549$  emission line, indicating the presence of outflows. More details about the xA sources behavior can be found in Martínez-Aldama et al. of this volume.

### 1.1. HE0359-3959: An Extreme xA Source

In our extreme luminosity Hamburg-ESO sample (Marziani et al., 2009; Sulentic et al., 2017), we have identified four cases of highly radiating quasars that show an extreme behavior, i.e., a high Eddington ratio and a strong blue asymmetry [ $c(1/2) < -4,000$  km s<sup>-1</sup>; centroid at half intensity] in the  $CIV\lambda 1549$  profile (Sulentic et al., 2017). The most extreme case corresponds to the quasars HE0359-3959, with  $z = 1.5209$ ,

$\log(L_{bol}) = 47.6$  erg s<sup>-1</sup> and a  $R_{FeII} = 1.12$ . It is cataloged as an A3 source (see **Figure 1**).

The aim of this paper is to analyze the spectral behavior of an extreme xA source, the quasar HE0359-3959. We performed multicomponent fits in a wide spectral range: UV, optical and Near-Infrared (Section 2); which gives us information about the dynamics and the physical conditions of the broad line region (BLR) (Section 3). In Section 4, we summarize the main results of our work.

## 2. OBSERVATIONS, DATA REDUCTION, AND MULTICOMPONENT FITTING

### 2.1. Observations and Data Reduction

Ultraviolet (UV), optical, and Near-Infrared spectra were observed with the Very Large Telescope (VLT-ESO). Optical and Near-Infrared spectra were obtained with the Infrared Spectrometer And Array Camera (ISAAC; decommissioned in 2013) using a slit of 0.6". The near-infrared spectrum was observed in 2010 in the K band with a total exposure time of 1,120 s. The optical spectrum was observed in 2004 in the J band with a total exposure time of 3,600 s. For the ultraviolet spectrum we used the Focal Reducer and low dispersion Spectrograph (FORS1) and a slit of 1.0" with a total exposure time of 1,440 s. It was observed in 2008. The data reduction was done using the IRAF package. The procedures followed are explained in Marziani et al. (2009), Martínez-Aldama et al. (2015), and Sulentic et al. (2017).

### 2.2. Multicomponent Fits

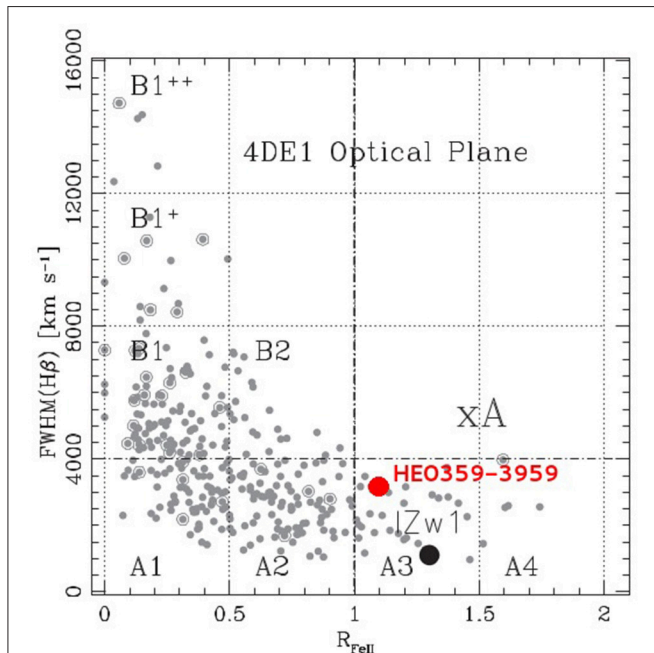
We perform multicomponent fits using SPECFIT, an IRAF routine (Kriss, 1994) to get the information of the most important emission lines. In each spectral range we fit a local continuum. The FWHM of all the broad components (BC) for  $H\beta$ ,  $AlIII\lambda 1860$ ,  $SiIII\lambda 1892$ ,  $CIV\lambda 1549$ , and  $SiIV\lambda 1397$  was taken equal. In the **Figure 2**, we present the multicomponents fits after continuum subtraction, for the  $CIV\lambda 1549$  and Ca II triplet range. The rest of the fits will be shown in an upcoming paper.

## 3. RESULTS

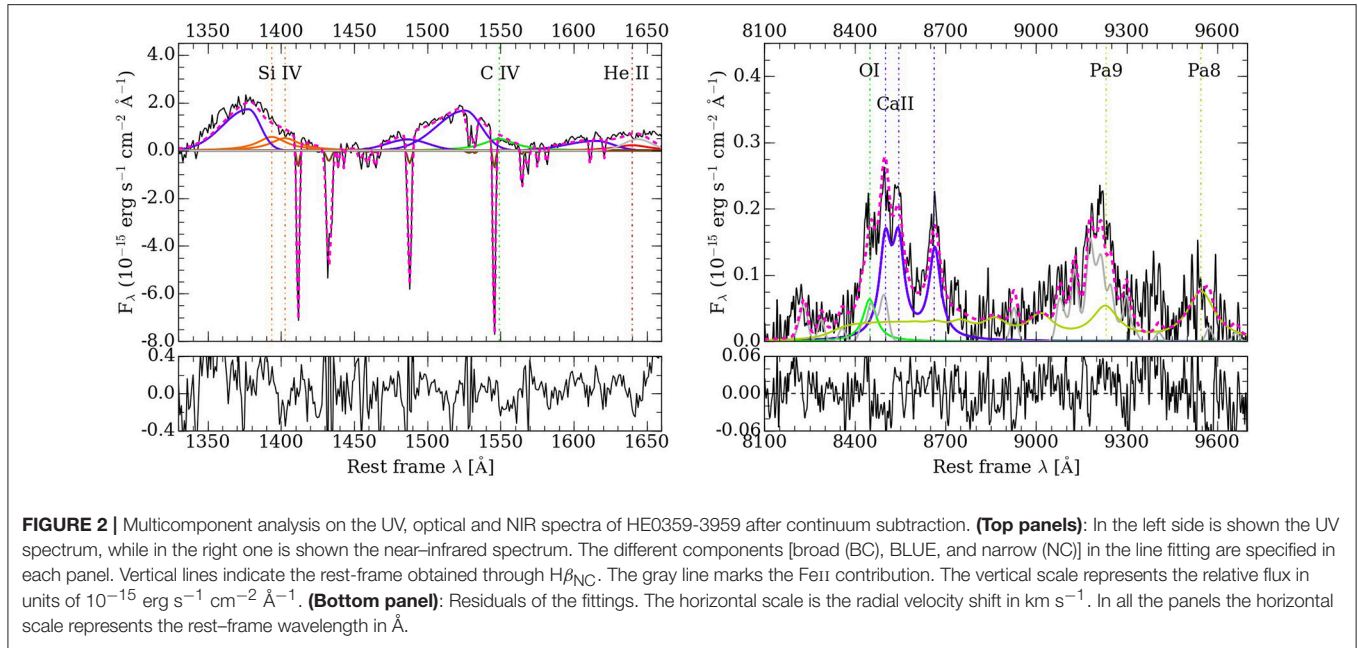
### 3.1. Multiwavelength Analysis

Low-ionization lines (LIL) have an ionization potential (IP)  $\lesssim 20$  eV. The  $H\beta$  line is the prototype of LIL. In population A3 and A4 sources  $H\beta$  has associated a blueshifted component (Bachev et al., 2004). In the case of HE0359-3959, the blueshifted component has a contribution to the total flux of  $\sim 9\%$ , and shows a centroid a half intensity of  $c(1/2) \approx -500 \pm 70$  km s<sup>-1</sup>.

The FeII (IP  $\sim 16$  eV) has an important contribution in the optical and near-infrared regions. To reproduce it we used the templates modeled by Marziani et al. (2009) and Garcia-Rissmann et al. (2012) for the optical and near-infrared ranges, respectively. Several works have found (Persson, 1988; Ferland and Persson, 1989; Joly, 1989; Dultzin-Hacyan et al., 1999; Martínez-Aldama et al., 2015) a close relationship between the FeII and the NIR Ca II  $\lambda 8498$ ,  $\lambda 8542$ , and  $\lambda 8662$  Å triplet. This relation is very well appreciable in this object: as well as the optical



**FIGURE 1** | 4DE1 Optical Plane reproduced from Marziani and Sulentic (2014). Gray points correspond to the sample of 470 bright low- $z$  QSOs from Zamfir et al. (2010). The plane is divided in bins according to Sulentic et al. (2002). Extreme accretor population A sources (xA) are located in A3 and A4 bins. The black dot indicates the position of 1 Zw 1, the prototype of low- $z$  xA sources. And, the red dot marks the location of HE0359-3959, an extreme xA source with high- $z$ .



Fe II is strong, the NIR Ca II triplet also is. It is the first time where we observe the Ca II triplet lines isolated at high redshift. Strong intensities of both ions imply an extremely low-ionization degree ( $U < 10^{-2}$ ;  $U$ : ionization parameter) and a high density ( $n_H \sim 10^{11-13} \text{ cm}^{-3}$ ) (Baldwin et al., 2004; Matsuoka et al., 2007; Martínez-Aldama et al., 2015).

In the UV region, the 1900 Å blend is formed by two intermediate-ionization lines (IIL;  $IP \sim 20\text{--}40 \text{ eV}$ ), Al III  $\lambda 1860$  and Si III  $\lambda 1892$ , which are accompanied by C III  $\lambda 1909$  and some Fe III transitions. In this blend we appreciate a blueshifted component. This component should be most likely associated with Al III  $\lambda 1860$ . Respect to Al III  $\lambda 1860$ , the blueshifted component has a contribution of the total profile of 60%. The centroid a half intensity is  $c(1/2) \approx -3,200 \pm 250 \text{ km s}^{-1}$ , which indicates the presence of an outflow generated by radiation forces presented in the intermediate-ionization lines (Marziani et al., 2017). On the other hand, considering the high intensity of Al III  $\lambda 1860$ , Si III  $\lambda 1892$ , Ca II and Fe II, it could suggest a possible chemical enrichment of the BLR (Juarez et al., 2009).

High ionization lines (HIL;  $IP > 40 \text{ eV}$ ), C IV  $\lambda 1549$ , He II  $\lambda 1640$ , and Si IV  $\lambda 1397$ , show a prominent blueshifted component. We find that the blue component has a contribution of 76, 62, and 57% to the total flux of C IV  $\lambda 1549$ , He II  $\lambda 1640$ , and Si IV  $\lambda 1397$ , respectively. The C IV  $\lambda 1549$  reaches  $c(1/2) \sim -6,000 \pm 500 \text{ km s}^{-1}$ , while He II  $\lambda 1640$  and Si IV  $\lambda 1397$   $c(1/2) \sim -4,000 \pm 550 \text{ km s}^{-1}$ . The velocities reached are ones of the highest found in the literature (Richards et al., 2011; Coatman et al., 2016; Sulentic et al., 2017). Then, it indicates that the full profile is dominated by an outflow and suggests the disk plus wind scenario (Gaskell, 1982; Richards et al., 2002, 2011).

### 3.2. Physical Properties of HE0359-3959

In order to study the physical properties of the quasar HE0359-3959, we built a grid of photoionization simulations using the

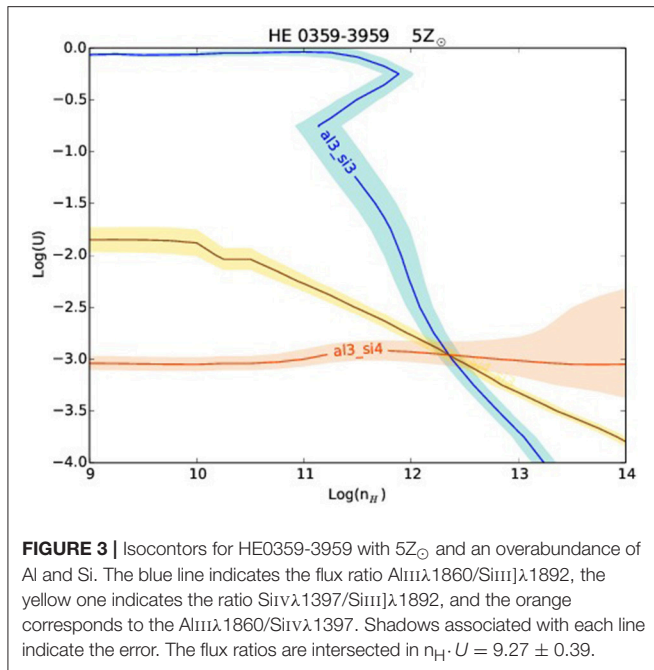
CLOUDY code (Ferland et al., 1998, 2013). For our simulations we consider a Mathews and Ferland continuum (Mathews and Ferland, 1987), a plane-parallel geometry, a metallicity  $5Z_{\odot}$  with an overabundance of Al and Si with respect to carbon (by a factor of three), and a column density of  $N_c = 10^{23} \text{ cm}^{-2}$ . See Negrete et al. (2012) for more details. Our simulations span the density range  $7.00 \leq \log(n_H) \leq 14.00$  and  $-4.5 \leq \log(U) \leq 0.00$  for the ionization parameter, in intervals of 0.25 dex. More details about the CLOUDY simulations can be found in Negrete et al. (2014). Using the UV lines, we define three groups of diagnostic ratios:

- The flux ratio Al III  $\lambda 1860$ /Si III  $\lambda 1892$  is a useful density diagnostic.
- The flux ratio Si IV  $\lambda 1397$ /Si III  $\lambda 1892$  for the ionization parameter.
- The flux ratio C IV  $\lambda 1549$ /Si IV  $\lambda 1397$  is mainly sensitive to the relative abundances of C and Si.

In Figure 3 is shown the result of the simulations. We obtained that the flux ratios are intersected in  $\log(n_H) = 12.32 \text{ cm}^{-3}$  and  $\log(U) = -2.95$ . Compared to not highly radiating AGNs (Negrete et al., 2013), this source shows a high density and a low ionization parameter, which marks a different behavior in the BLR, probably caused by the slim disk hosted in these kind of sources. Taking into account the high intensity of Al III  $\lambda 1860$ , Fe II, and Ca II we conclude that effectively the low-ionization emitter zone has a high density and low-ionization parameter.

Negrete et al. (2012) proposed a new method to determine the size of the BLR ( $r_{\text{BLR}}$ ) and the black hole mass ( $M_{\text{BH}}$ ) based on the product  $n_H \cdot U$  and independently of redshift. This method gives similar results to the obtained from the classical methods such as reverberation mapping at low- $z$  (Negrete et al., 2014). Knowing the product of  $n_H \cdot U$  obtained from the CLOUDY simulations, we compute the size of the BLR ( $r_{\text{BLR}}$ ) and considering the





FWHM of the broad components as the velocity dispersion, we can get the black hole mass ( $M_{\text{BH}}$ ) and the Eddington ratio. The size of the BLR is  $\log(r_{\text{BLR}}) = 18.37 \pm 0.04$  cm and the black hole mass is  $\log(M_{\text{BH}}) = 9.52 \pm 0.41 M_{\odot}$ . These values are in agreement with the ones found for a large xA sample at high-redshift (Martínez-Aldama et al., in preparation).

The Eddington ratio for this source is  $L/L_{\text{Edd}} = 0.74 \pm 0.11$ . Considering that it shows a  $c(1/2) \sim -6,000 \pm 500 \text{ km s}^{-1}$  for  $\text{C IV } \lambda 1549$ , we confirm the directly proportional relation between  $c(1/2)$  and  $L/L_{\text{Edd}}$ . Indicating that  $L/L_{\text{Edd}}$  could be the driver of the outflows (Sulentic et al., 2017).

## REFERENCES

- Abramowicz, M. A., Czerny, B., Lasota, J. P., and Szuszkiewicz, E. (1988). Slim accretion disks. *Astrophys. J.* 332, 646–658.
- Abramowicz, M. A., and Straub, O. (2014). Accretion discs. *Scholarpedia* 9:2408. doi: 10.4249/scholarpedia.2408
- Bachev, R., Marziani, P., Sulentic, J. W., Zamanov, R., Calvani, M., and Dultzin-Hacyan, D. (2004). Average ultraviolet quasar spectra in the context of eigenvector 1: a baldwin effect governed by the eddington ratio? *Astrophys. J.* 617, 171–183. doi: 10.1086/425210
- Baldwin, J. A., Ferland, G. J., Korista, K. T., Hamann, F., and LaCluyzé, A. (2004). The origin of Fe II emission in active galactic nuclei. *Astrophys. J.* 615, 610–624. doi: 10.1086/424683
- Coatman, L., Hewett, P. C., Banerji, M., and Richards, G. T. (2016). C IV emission-line properties and systematic trends in quasar black hole mass estimates. *Mon. Not. R. Astron. Soc.* 461, 647–665. doi: 10.1093/mnras/stw1360
- Dultzin-Hacyan, D., Taniguchi, Y., and Uranga, L. (1999). “Where is the Ca II triplet emitting region in AGN?,” in *Structure and Kinematics of Quasar Broad Line Regions*, Vol. 175, *Astronomical Society of the Pacific Conference Series*, eds C. M. Gaskell, W. N. Brandt, M. Dietrich, D. Dultzin-Hacyan, and M. Eracleous (San Francisco, CA: Astronomical Society of the Pacific), 303.

## 4. CONCLUSIONS

The information given by the multiwavelength analysis indicates that in HE0359-3959 there is coexistence of substructures in the broad line region. Low and intermediate-ionization regions, where  $\text{H}\beta$ ,  $\text{Al III } \lambda 1860$  and  $\text{Si III } \lambda 1892$  are emitted, are dense ( $n_{\text{H}} \sim 10^{11-12} \text{ cm}^{-3}$ ) and optically thick ( $U \sim 10^{-2.5}$ ). They are mainly governed by virial motions and the presence of a blueshifted component indicates the influence of radiation forces. On the other hand, according to Marziani et al. (2010) the high-ionization region is less dense ( $n_{\text{H}} \sim 10^{10} \text{ cm}^{-3}$ ,  $U \sim 10^{-1}$ ), pointing out a difference with the physical conditions shown by the low and intermediate-ionization lines.

High ionization lines are dominated by strong radiation forces, producing outflows in high-ionization lines like  $\text{C IV } \lambda 1549$ ,  $\text{He II } \lambda 1640$ , and  $\text{Si IV } \lambda 1397$ . The high Eddington ratio value suggests the presence of a slim optically thick disk which could be related to the extreme outflow properties observed in HE0359-3959. The presence of strong outflows has been related with the co-evolution of the active galactic nuclei and the host galaxy.

## AUTHOR CONTRIBUTIONS

MLM-A and PM: Data reduction, multicomponents fits, analysis, writing, revision. AD: Analysis, reduction, writing, revision. CN: Analysis, photoionization models, revision. DD: Analysis, revision. MAM-C: Data reduction, multicomponents fits.

## ACKNOWLEDGMENTS

MLM-A acknowledge the postdoctoral grant from the CONACyT. MLM-A, AD, and MAM-C acknowledge financial support from Spanish Ministry for Economy and Competitiveness through grants AYA2013-42227-P and AYA2016-76682-C3-3-1-P.

- Ferland, G. J., Korista, K. T., Verner, D. A., Ferguson, J. W., Kingdon, J. B., and Verner, E. M. (1998). CLOUDY 90: numerical simulation of plasmas and their spectra. *Publ. Astron. Soc. Pac.* 110, 761–778.
- Ferland, G. J., and Persson, S. E. (1989). Implications of CA II emission for physical conditions in the broad-line region of active galactic nuclei. *Astrophys. J.* 347, 656–673.
- Ferland, G. J., Porter, R. L., van Hoof, P. A. M., Williams, R. J. R., Abel, N. P., Lykins, M. L., et al. (2013). The 2013 release of cloudy. *RMxAA* 49, 137–163.
- García-Rissmann, A., Rodríguez-Ardila, A., Sigut, T. A. A., and Pradhan, A. K. (2012). A near-infrared template derived from I Zw 1 for the Fe II emission in active galaxies. *Astrophys. J.* 751:7. doi: 10.1088/0004-637X/751/1/7
- Gaskell, C. M. (1982). A redshift difference between high and low ionization emission-line regions in QSOs - Evidence for radial motions. *Astrophys. J.* 263, 79–86.
- Joly, M. (1989). Formation of CA II lines in active galactic nuclei. *Astron. Astrophys.* 208, 47–51.
- Juarez, Y., Maiolino, R., Mújica, R., Pedani, M., Marinoni, S., Nagao, T., et al. (2009). The metallicity of the most distant quasars. *Astron. Astrophys.* 494, L25–L28.
- Kriss, G. (1994). “Fitting models to UV and optical spectral data,” in *Astronomical Data Analysis Software and Systems III*, Vol. 61 *Astronomical Society of the*

- Pacific Conference Series*, eds D. R. Crabtree, R. J. Hanisch, and J. Barnes (San Francisco, CA: Astronomical Society of the Pacific), 437.
- Martínez-Aldama, M. L., Dultzin, D., Marziani, P., Sulentic, J. W., Bressan, A., Chen, Y., et al. (2015). O I and Ca II observations in intermediate redshift quasars. *Astrophys. J.* 217:3. doi: 10.1088/0067-0049/217/1/3
- Marziani, P., Del Olmo, A., Martínez-Aldama, M. L., Dultzin, D., Negrete, C., Bon, E., et al. (2017). Quasar black hole mass estimates from high-ionization lines: breaking a taboo? *Atoms* 5, 33–47. doi: 10.3390/atoms5030033
- Marziani, P., and Sulentic, J. W. (2014). Highly accreting quasars: sample definition and possible cosmological implications. *Mon. Not. R. Astron. Soc.* 442, 1211–1229. doi: 10.1093/mnras/stu951
- Marziani, P., Sulentic, J. W., Negrete, C. A., Dultzin, D., Zamfir, S., and Bachev, R. (2010). Broad-line region physical conditions along the quasar eigenvector 1 sequence. *Mon. Not. R. Astron. Soc.* 409, 1033–1048. doi: 10.1111/j.1365-2966.2010.17357.x
- Marziani, P., Sulentic, J. W., Stirpe, G. M., Zamfir, S., and Calvani, M. (2009). VLT/ISAAC spectra of the H $\beta$  region in intermediate-redshift quasars. III. H $\beta$  broad-line profile analysis and inferences about BLR structure. *Astron. Astrophys.* 495, 83–112. doi: 10.1051/0004-6361/200810764
- Mathews, W. G., and Ferland, G. J. (1987). What heats the hot phase in active nuclei? *Astrophys. J.* 323, 456–467.
- Matsuoka, Y., Oyabu, S., Tsuzuki, Y., and Kawara, K. (2007). Observations of O I and Ca II emission lines in Quasars: implications for the site of Fe II line emission. *Astrophys. J.* 663, 781–798. doi: 10.1086/518399
- Negrete, C. A., Dultzin, D., Marziani, P., and Sulentic, J. W. (2012). Broad-line region physical conditions in extreme population a Quasars: a method to estimate central black hole mass at high redshift. *Astrophys. J.* 757:62. doi: 10.1088/0004-637X/757/1/62
- Negrete, C. A., Dultzin, D., Marziani, P., and Sulentic, J. W. (2013). Reverberation and photoionization estimates of the broad-line region radius in low-z Quasars. *Astrophys. J.* 771:31. doi: 10.1088/0004-637X/771/1/31
- Negrete, C. A., Dultzin, D., Marziani, P., and Sulentic, J. W. (2014). A new method to obtain the broad line region size of high redshift Quasars. *Astrophys. J.* 794:95. doi: 10.1088/0004-637X/794/1/95
- Persson, S. E. (1988). Calcium infrared triplet emission in active galactic nuclei. *Astrophys. J.* 330, 751–765.
- Richards, G. T., Kruczek, N. E., Gallagher, S. C., Hall, P. B., Hewett, P. C., Leighly, K. M., et al. (2011). Unification of luminous type 1 Quasars through C IV emission. *Astrophys. J.* 141, 167–183. doi: 10.1088/0004-6256/141/5/167
- Richards, G. T., Vanden Berk, D. E., Reichard, T. A., Hall, P. B., Schneider, D. P., SubbaRao, M., et al. (2002). Broad emission-line shifts in Quasars: an orientation measure for radio-quiet quasars? *Astrophys. J.* 124, 1–17. doi: 10.1086/341167
- Sulentic, J. W., Del Olmo, A., Marziani, P., Martínez-Carballo, M. A., D'Onofrio, M., Oyabu, S., et al. (2017). What does CIV1549 tell us about the physical driver of the Eigenvector Quasar Sequence? *Astron. Astrophys. arXiv:1708.03187*
- Sulentic, J. W., Marziani, P., and Dultzin-Hacyan, D. (2000a). Phenomenology of broad emission lines in active galactic nuclei. *Annu. Rev. Astron. Astrophys.* 38, 521–571. doi: 10.1146/annurev.astro.38.1.521
- Sulentic, J. W., Marziani, P., Zamanov, R., Bachev, R., Calvani, M., and Dultzin-Hacyan, D. (2002). Average quasar spectra in the context of eigenvector 1. *Astrophys. J.* 566, L71–L75. doi: 10.1086/339594
- Sulentic, J. W., Zwitter, T., Marziani, P., and Dultzin-Hacyan, D. (2000b). Eigenvector 1: an optimal correlation space for active galactic nuclei. *Astrophys. J.* 536, L5–L9.
- Zamfir, S., Sulentic, J. W., Marziani, P., and Dultzin, D. (2010). Detailed characterization of H $\beta$  emission line profile in low-z SDSS quasars. *Mon. Not. R. Astron. Soc.* 403, 1759–1786. doi: 10.1111/j.1365-2966.2009.16236.x

**Conflict of Interest Statement:** The authors declare that the research was conducted in the absence of any commercial or financial relationships that could be construed as a potential conflict of interest.

Copyright © 2017 Martínez-Aldama, Del Olmo, Marziani, Negrete, Dultzin and Martínez-Carballo. This is an open-access article distributed under the terms of the Creative Commons Attribution License (CC BY). The use, distribution or reproduction in other forums is permitted, provided the original author(s) or licensor are credited and that the original publication in this journal is cited, in accordance with accepted academic practice. No use, distribution or reproduction is permitted which does not comply with these terms.